

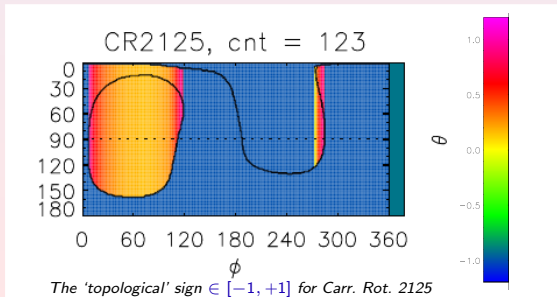
# Implications of Solar Magnetograms for the Drifts of Cosmic Rays

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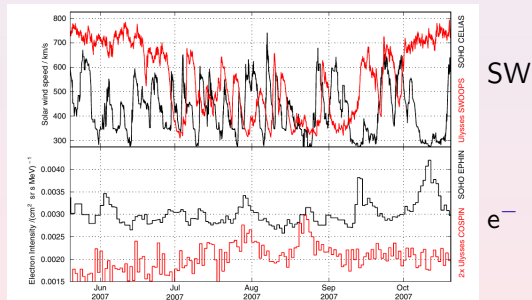


# Motivation

A **guiding question** for contemporary research activities:

*Can a combination of current models of the kinetic transport of cosmic rays with magnetohydrodynamic (MHD) models of the solar wind and its turbulence be used to explain three-dimensional (3D) multi-point spacecraft data?*

Example: Solar wind and energetic electron data at SOHO (black) and Ulysses (red)



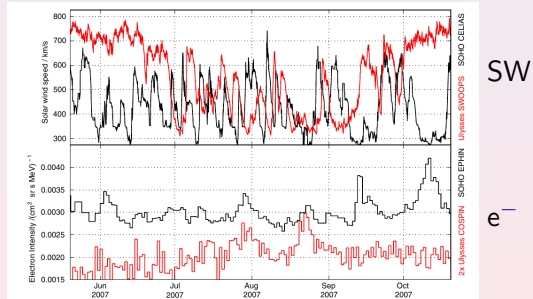
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→ *Task: Construction of suitable combination of models*

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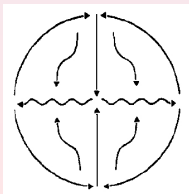
- Combination of MHD solar wind and kinetic cosmic ray transport models has long tradition  
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(e.g., ENGELBRECHT & BURGER [2013], WIENGARTEN ET AL. [2016], MOLOTO ET AL. [2018])
- Particularly, a physics-based modelling of (the reduction of) drifts in a structured and time-dependent solar wind that is fully consistent with observations is still missing  
(e.g., MOLOTO ET AL. [2018], ZHAO ET AL. [2018], KOPP ET AL. [2021])



$$qA > 0$$

# Outline of the Talk

- Introduction (✓)
- 'Traditional' models of drift reduction

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- Résumé

# Traditional models of drift reduction

Heliospheric particle drifts are described by

$$\langle \mathbf{v}_d \rangle = \nabla \times \left( \kappa_A \frac{\mathbf{B}}{B} \right) ; \kappa_A = qA\kappa_{A,0} \frac{P\beta}{3B} \frac{(P/P_0)^2}{1 + (P/P_0)^2} ; \kappa_{A,0} \in [0, 1]$$

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## *Reduction Method I:*

Phenomenological coupling to  
(current sheet) tilt angle  $\alpha$ :

$$\kappa_{A,0} = \left[ \cos \left( \frac{\pi}{150^\circ} \alpha \right) \right]^{\frac{\alpha}{c_1}}$$

FERREIRA & POTGIETER [2003],  
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Direct coupling  
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Problems: parametrization with  $\alpha$  too simple;  $\delta B/B$  exhibits  
only weak, if any, dependence on solar cycle

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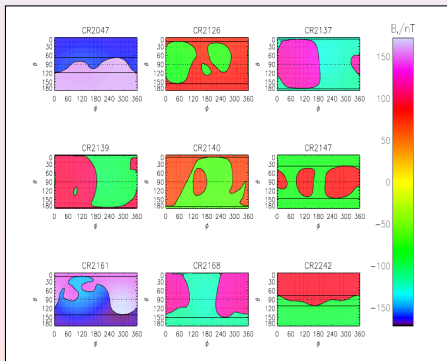
## *New Reduction Method:*

Coupling to topology of magnetic field maps (see figure):

$$A \rightarrow \sigma_t \hat{=} \text{'topological' sign}$$

KOPP ET AL. [2021]

- $A = \pm 1$  discrete
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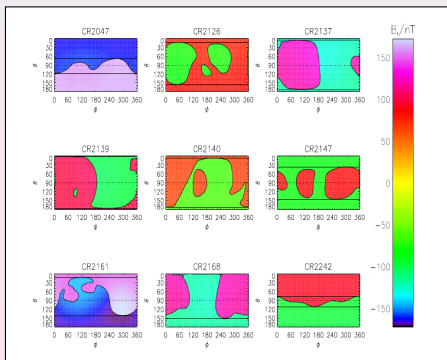
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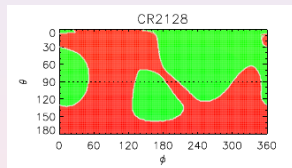
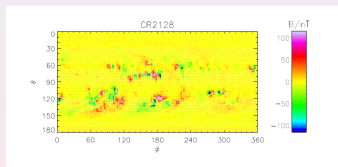
- $A = \pm 1$  discrete
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Advantage: allows for localized regions with ordered drift motion and is using the MHD boundary conditions

# Computation of the topological sign

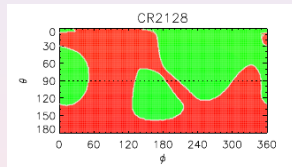
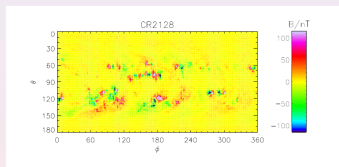
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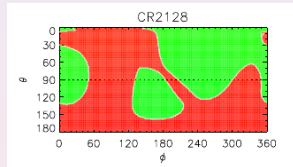
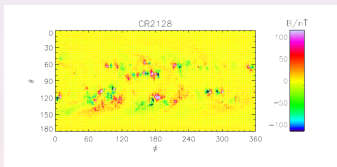
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  - new tilt angles can be greater than Wilcox tilt angles which are limited to  $\pm 75$  deg

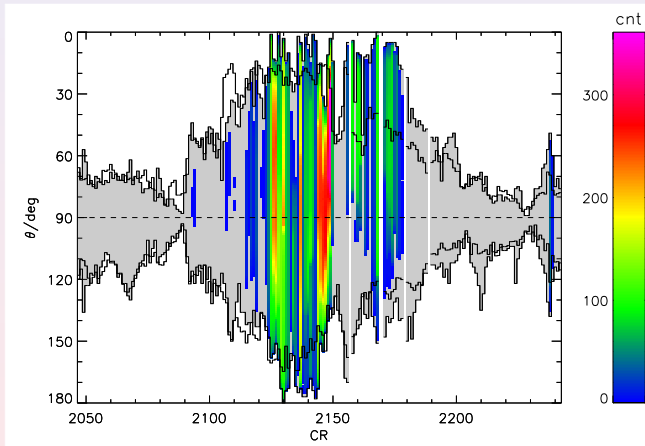
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- **Step 2:** Computation of topological sign  $\sigma_t \in [-1, +1]$  that takes into account localized regions of opposite field polarities

# Step 1: New vs. Wilcox tilt angles



- new angles: gray area, thick lines
- WSO angles: thin lines
- colored bars: extent of region with multiple HCS crossings

## Step 2: Topological sign

Distinguish sign of 'inner' and 'outer' regions (KOPP ET AL. [2021]):

$$\sigma_t = \langle \sigma_{t,\text{out}}(\varphi) + w(\varphi)\sigma_{t,\text{inn}}(\varphi) \rangle_\varphi; \quad w(\varphi) = (\tilde{\nu}_{\text{max}}(\varphi) - \tilde{\nu}_{\text{min}}(\varphi))/180^\circ$$

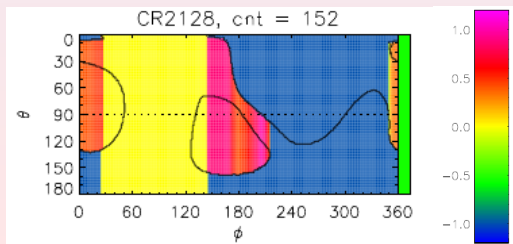
with

$$\sigma_{t,\text{out}}(\varphi) = \frac{1}{2} \left( q_r(\tilde{\nu}_{\text{min}}(\varphi) - 1, \varphi) - q_r(\tilde{\nu}_{\text{max}}(\varphi) + 1, \varphi) \right)$$

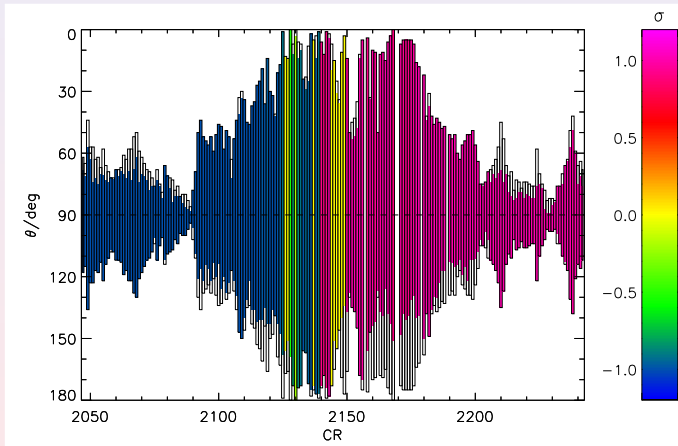
and

$$q_r(\vartheta, \varphi) = \frac{B_r(\vartheta, \varphi)}{|B_r(\vartheta, \varphi)|}$$

( $\sigma_{t,\text{inn}}$  similar, see ICRC paper)



## Step 2: Topological sign vs. Carrington Rotation



- new tilt angles with topological sign
- box extensions indicate mirrored larger value of other hemisphere
- sign of the cycle changes between CRs 2139 and 2140

# Résumé

## The **topological sign**

- *generalizes previous approaches, which were limited to  $A = \pm 1$ , to  $\sigma_t \in [-1, +1]$*
- *takes into account the topology of magnetic field maps at the heliobase and is, thus, explicitly physics-based (as opposed to heuristic)*
- *enables a consistent coupling of kinetic transport and MHD models in the sense that they are employing the same boundary conditions*
- *approach may be supplemented by additional effects due to, e.g., turbulence*