



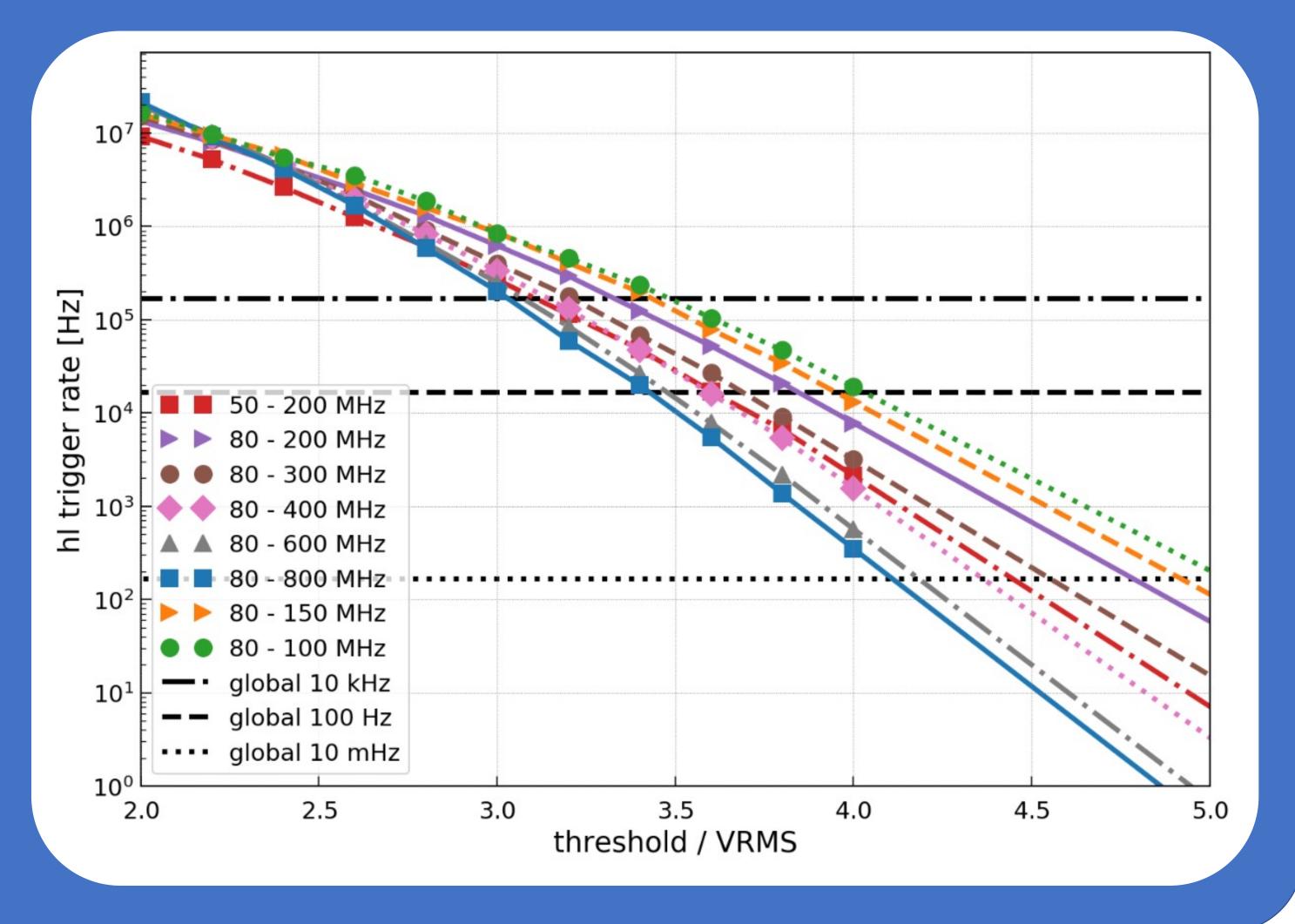
Ultra-high energy neutrino astronomy

- Measurement of EeV neutrinos next milestone in astroparticle physics
- Radio technology explored in test-bed arrays
- Now, several mid and large-scale arrays being constructed or planned
 - RNO-G (Greenland, under construction, #1058)
 - ARIANNA-200 (Ross Ice shelf, planned #1190)
 - IceCube-Gen2 (South Pole, planned #1183)
- Challenge: Low flux and small interaction cross section \rightarrow low statistics
- Solution: Optimize sensitivity of each detector station
- here: optimization of trigger bandwidth

Method to compare different trigger schemes

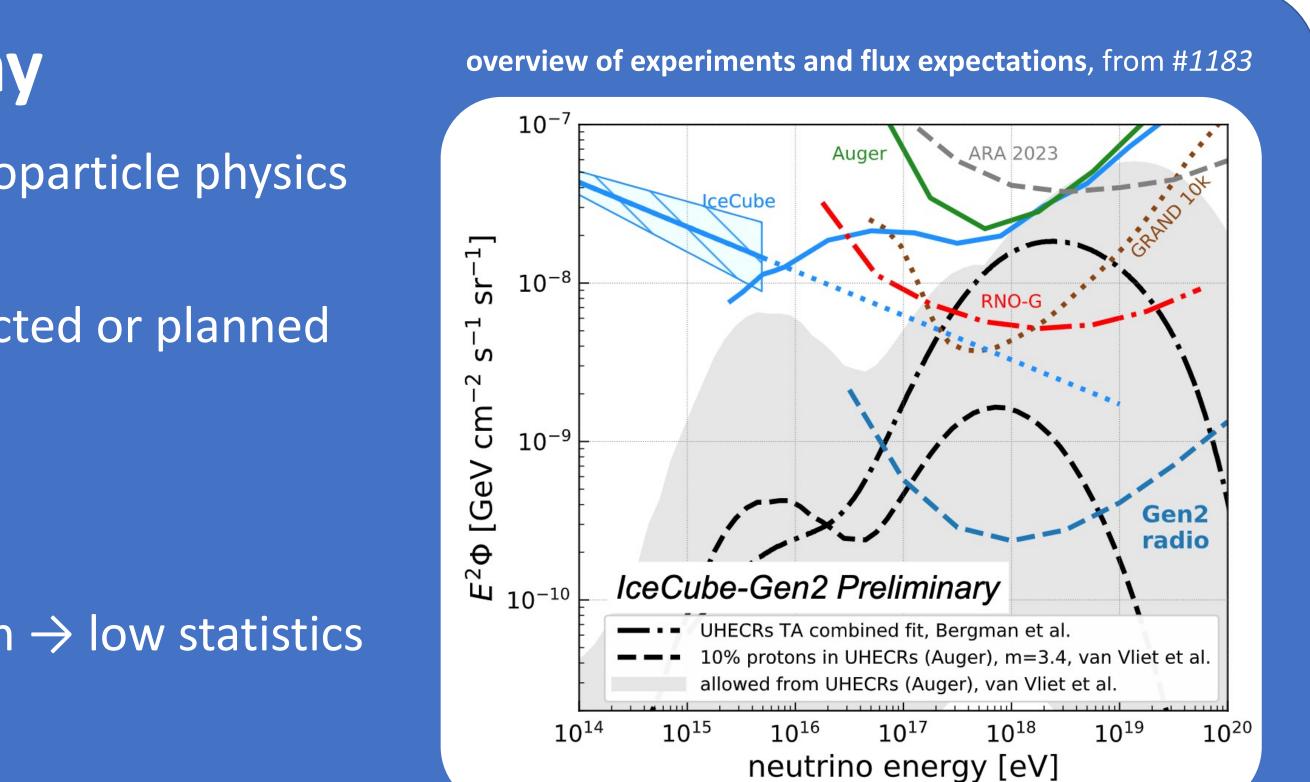
- Trigger rate dominated by thermal noise fluctuations \rightarrow max. data rate defines the threshold
- Correct comparisons of trigger schemes:
- Calculate trigger thresholds that yield the same noise trigger rate
- Studied here:
- high/low amplitude threshold trigger
- additional coincidence requirement 2 out of 4 antennas
- noise trigger rates depend on threshold and bandwidth
- Method can also be used to compare a power integration trigger with an amplitude threshold trigger

 \rightarrow see JINST 16 T05001 (2021) for details



Optimization of trigger bandwidth \rightarrow 50% increase in neutrino sensitivity for radio detectors

An improved trigger for Askaryan radio detectors

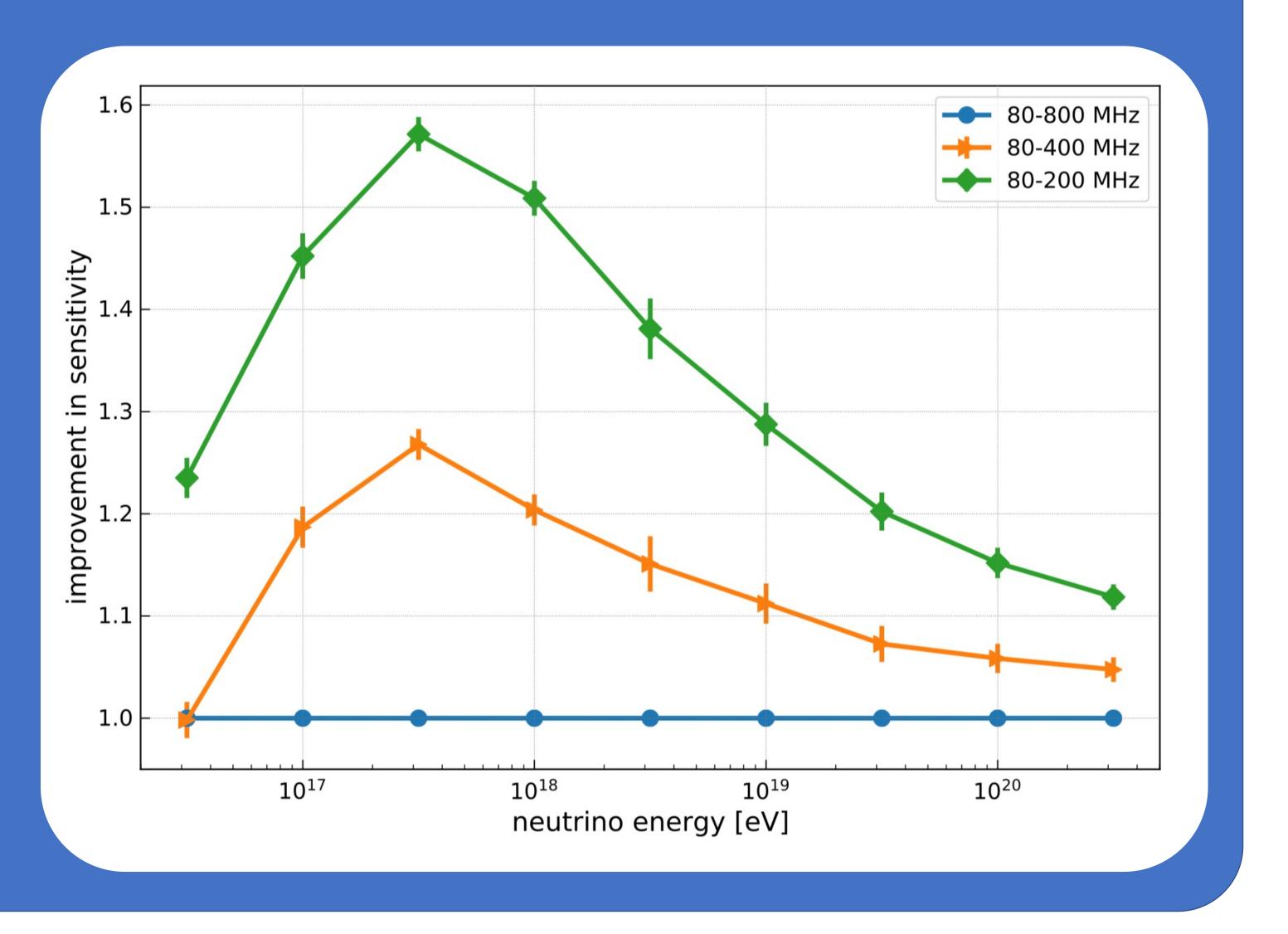


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Sensitivity improvement

- Current Askaryan detectors operate at large bandwidth of 80 MHz to ~1 GHz
- Studied here:
 - ARIANNA detector on Ross Ice Shelf
- Restricting trigger bandwidth to 80-200 MHz
- 50% improvement between 10¹⁷ and 10¹⁸eV



Reason for improvement

- 1. Off-cone events with low frequency content are frequent
- 2. Antennas more sensitive at low frequencies
 - sensitivity ~ 1/frequency
- 3. Noise RMS flat in frequency
 - \rightarrow noise can be reduced by reducing bandwidth

Average measured frequency spectrum shows clear peak at low frequencies

PoS(ICRC2021)1050

Averaged measured frequency spectrum of simulated neutrinos with an energy of 10 17eV

