Solar Atmospheric Neutrinos Searches with the ANTARES Neutrino Telescope

(Contribution #534)

ANTARES-KM3NeT
37th ICRC-2021

July 16th, 2021

University of Granada

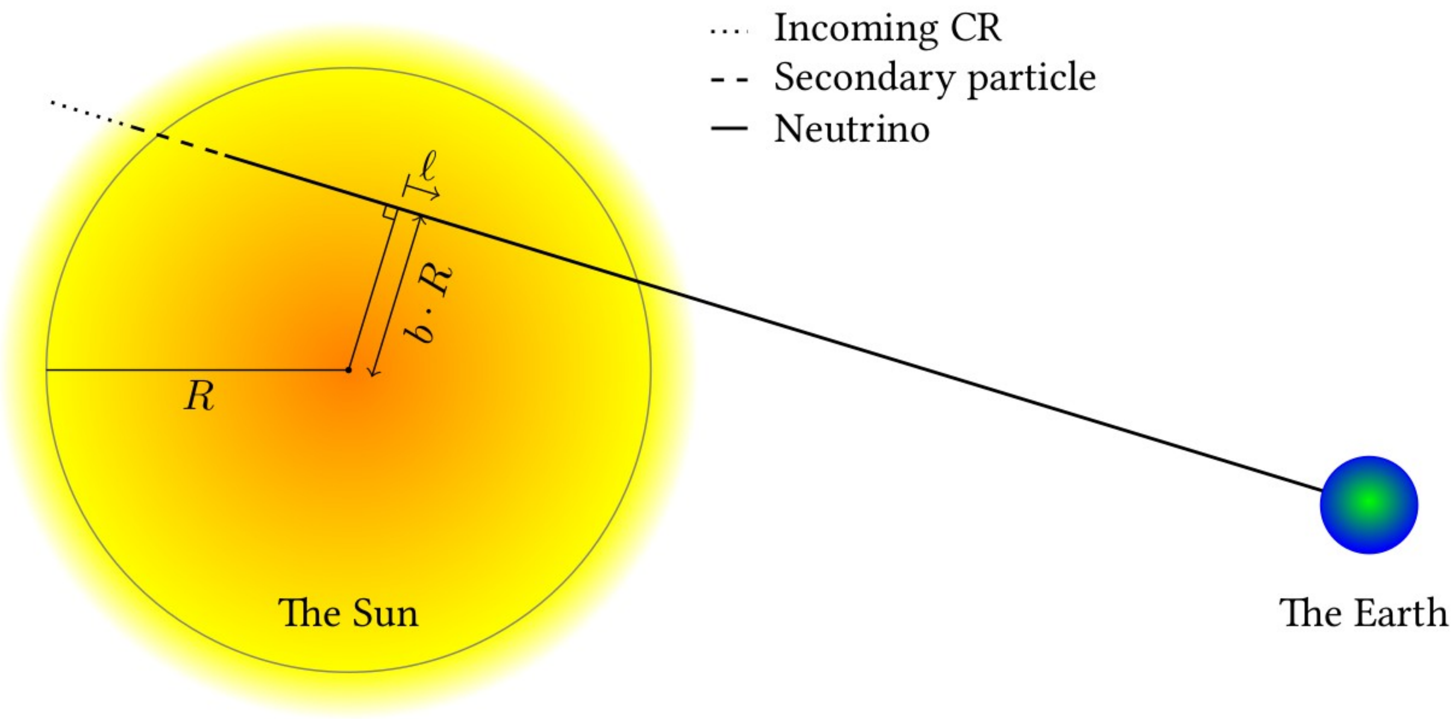
On behalf of the ANTARES collaboration

D. Lopez-Coto
S. Navas
J.D. Zornoza

Solar Atmospheric Neutrinos

- CRs blocked by the Sun yield ν as final state particles.
- The majority of the neutrinos are absorbed in the inner part.
- ν produced at the solar corona can escape and reach the Earth.
- Important for understanding the solar composition as well as the background for indirect solar DM searches.

Figure from:
J. Edsjö et al JCAP06(2017)033



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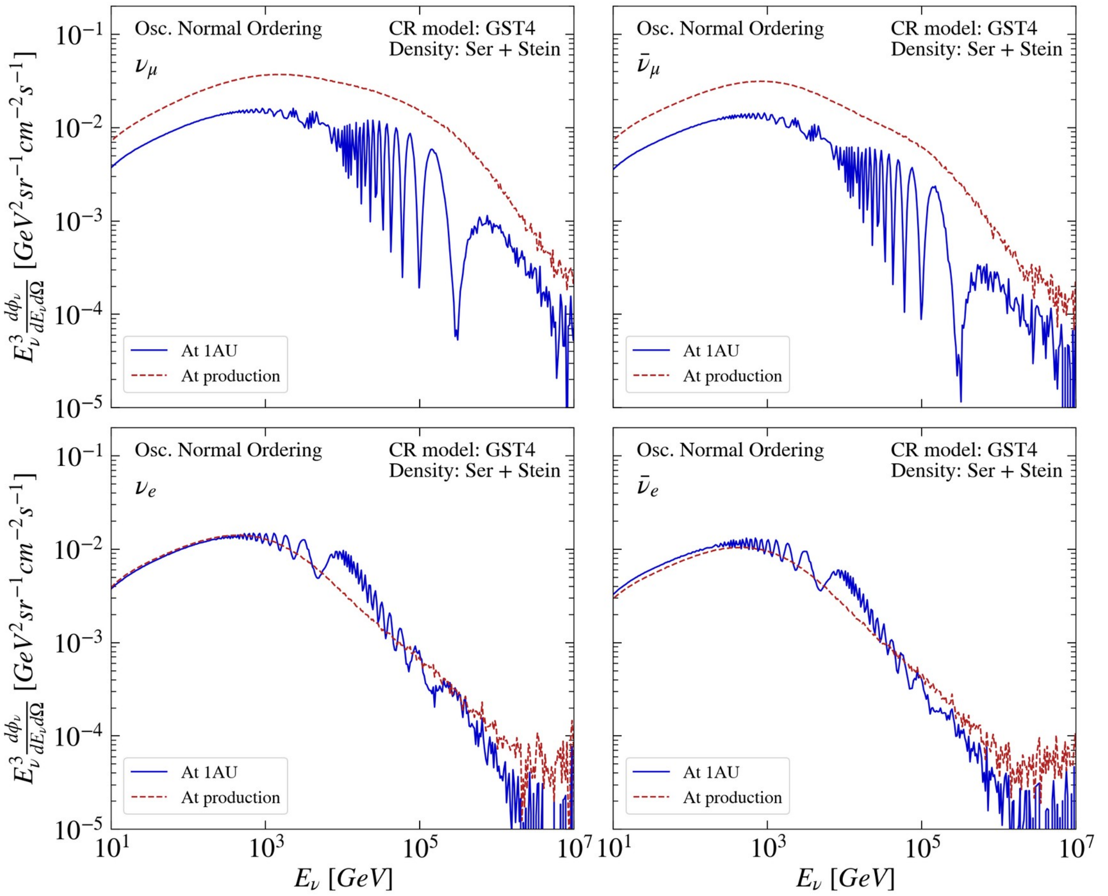
Solar Atmospheric Neutrinos

ANTARES – Solar Atmospheric Neutrinos

- Solar Atmospheric Neutrino flux from WimpSim
  From: J. Edsjö et al JCAP06(2017)033
  - 2 Cosmic Ray (CR) models (H3a and GST4).
  - 2 Solar composition models. (Ser+Stein and Ser+GS98).
  - Oscillation and Normal Ordering parameters. From global-best fit: JHEP 01 (2017) 087

$$\theta_{12} = 33.56^\circ \qquad \delta = 261^\circ$$
$$\theta_{13} = 8.46^\circ \qquad \Delta m^2_{21} = 7.5 \cdot 10^{-5}\mathrm{eV}^2$$
$$\theta_{23} = 41.6^\circ \qquad \Delta m^2_{31} = 2.524 \cdot 10^{-3}\mathrm{eV}^2$$

- Sun path for 2008-2018 period.



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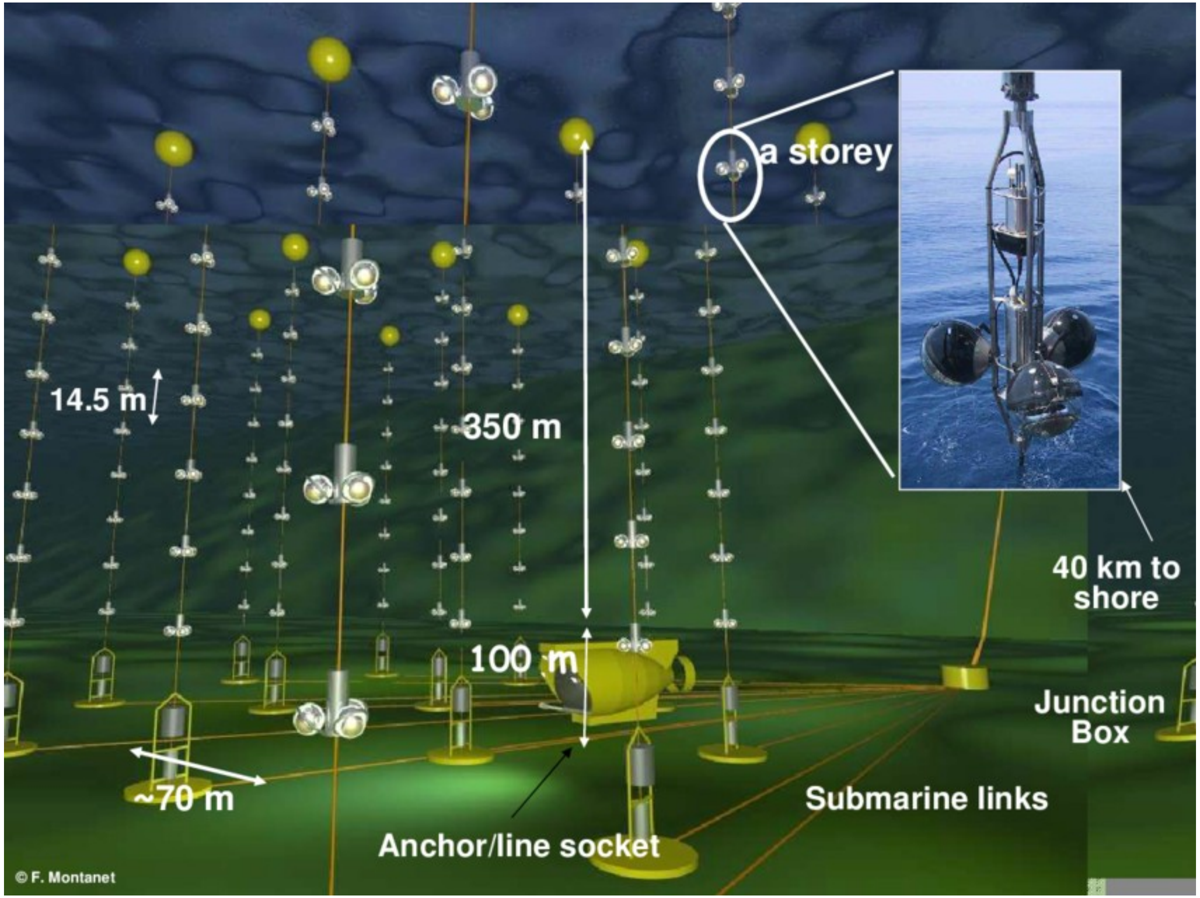
The ANTARES Detector

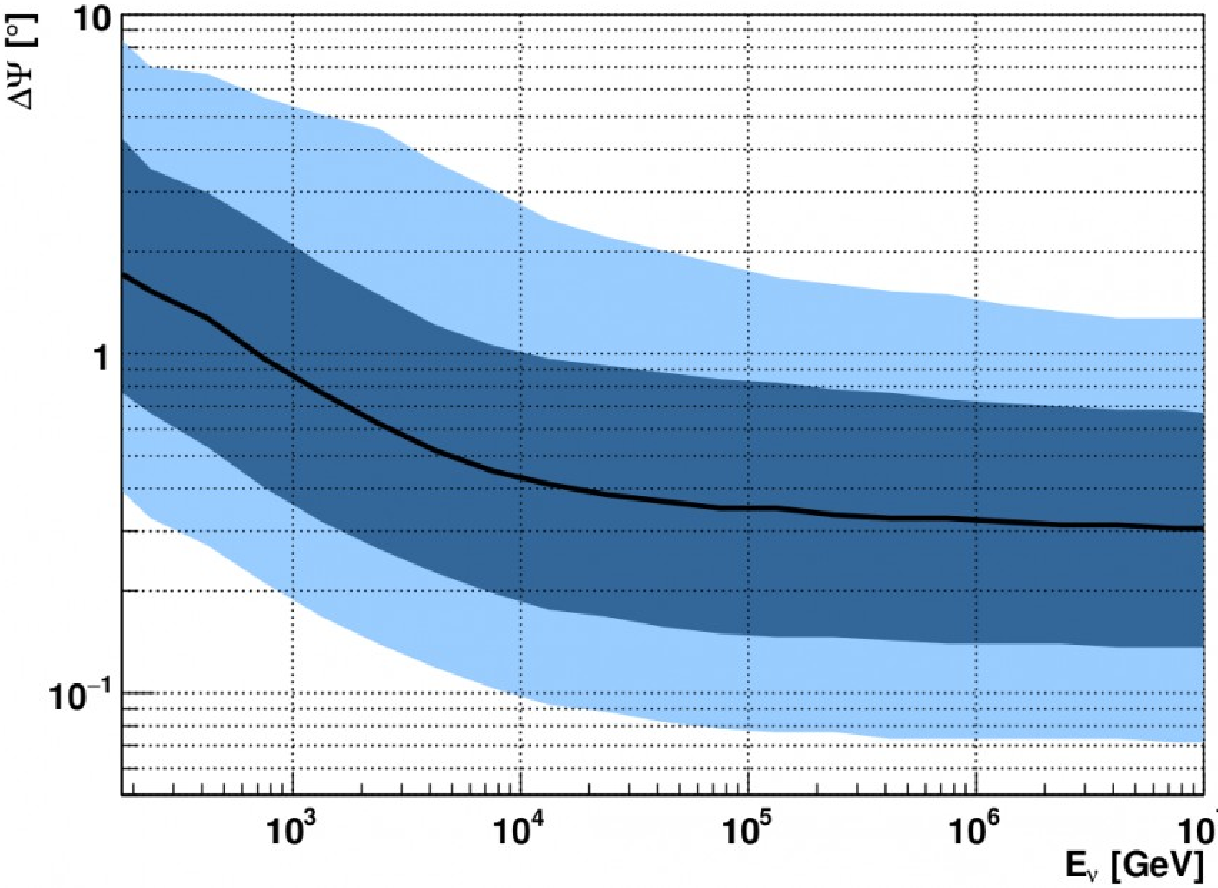
- First undersea Neutrino Telescope.
- Located in the Mediterranean Sea, near Toulon, at 2500 m depth.
- Construction 2006-2008.
- Continuously taking data.
- 12 lines (885 PMTs)
- 25 storeys/line
- 3 PMTs/storey



The ANTARES Detector

- 2 main topologies
  - Track like → From ν_μ and ν_τ CC.
  - Shower like→ From all-flavours NC and ν_e and ν_τ CC.
- Angular resolution < 0.4° for E_ν > 10 TeV).

Median angular resolution, track channel.



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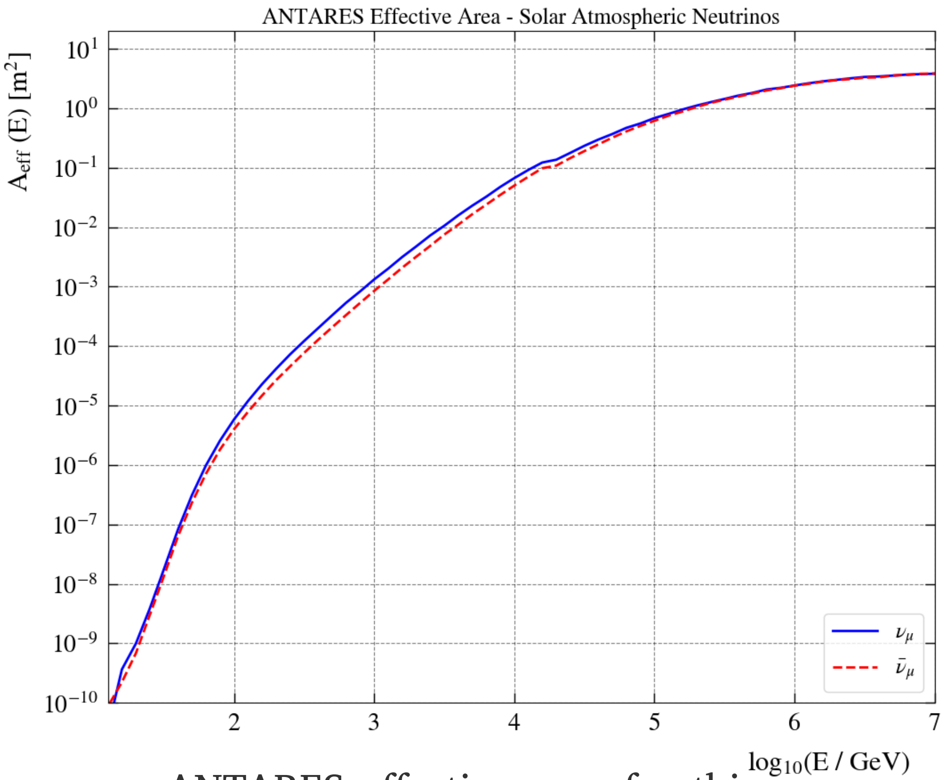
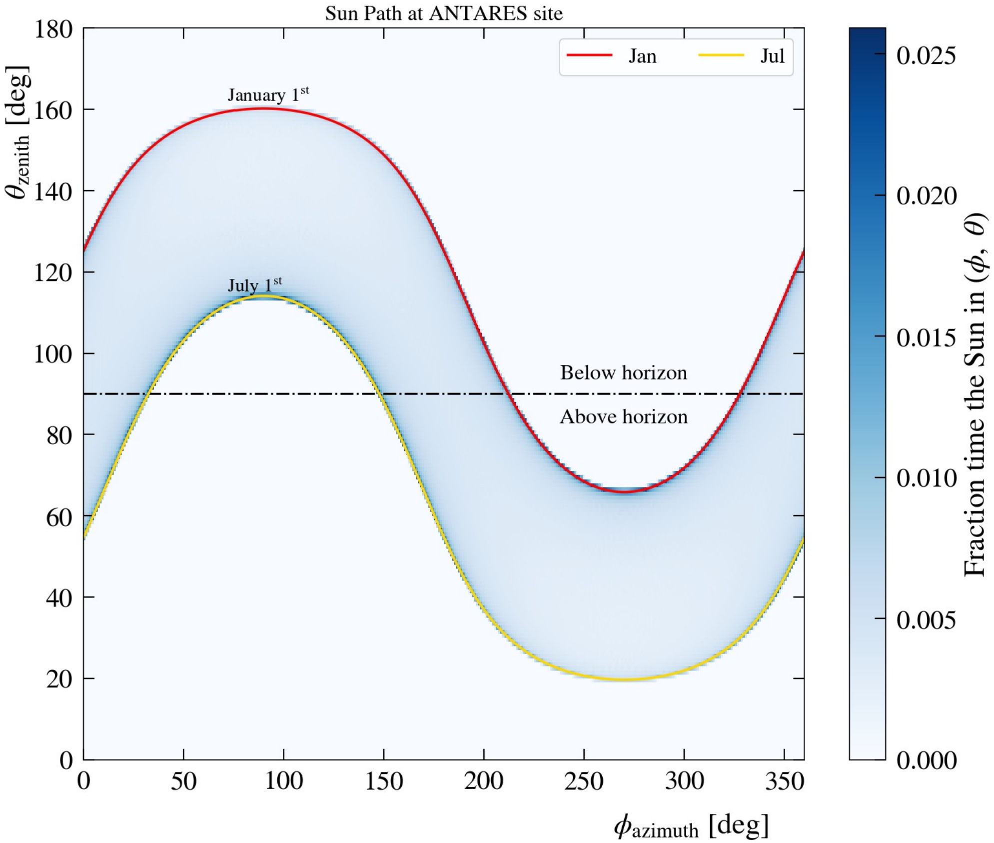
The ANTARES Detector





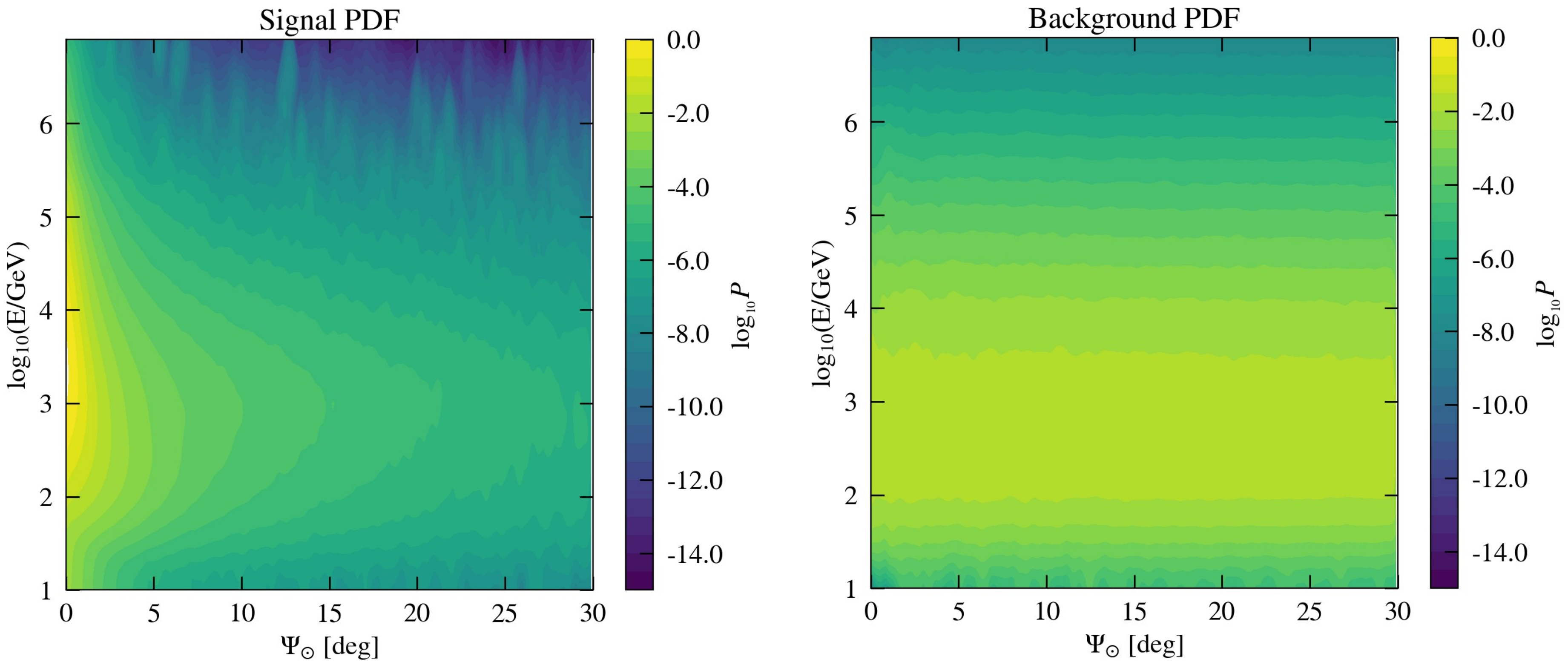
- ANTARES effective area for this analysis.
- Sun tracking taken into account.

Analysis

- Only track channel considered (ν_μ CC).
- Data taking period from 2008 to 2018 (both included) → lifetime of 3022 days.
- Main background → Atmospheric μ and atmospheric ν.
- Selection quality cuts to optimize Sensitivity and reject background.
  - $\Lambda > -5.2$,
  - $\beta < 1°$,
  - $\cos\theta > 0$ → upward-going events.
- Unbinned likelihood search.

Analysis

- Unbinned Likelihood search. $\mathcal{L}(n_{\rm sig}) = e^{-(n_{\rm sig}+n_{\rm bkg})} \prod_{i}^{N} \left[n_{\rm sig} \cdot \mathcal{S}(\Psi_{\odot,i}, \beta_i, E_i) + n_{\rm bkg} \cdot \mathcal{B}(\Psi_{\odot,i}, \beta_i, E_i) \right]$
- Signal and Background PDFs from MC weighted events and scrambled data respectively.



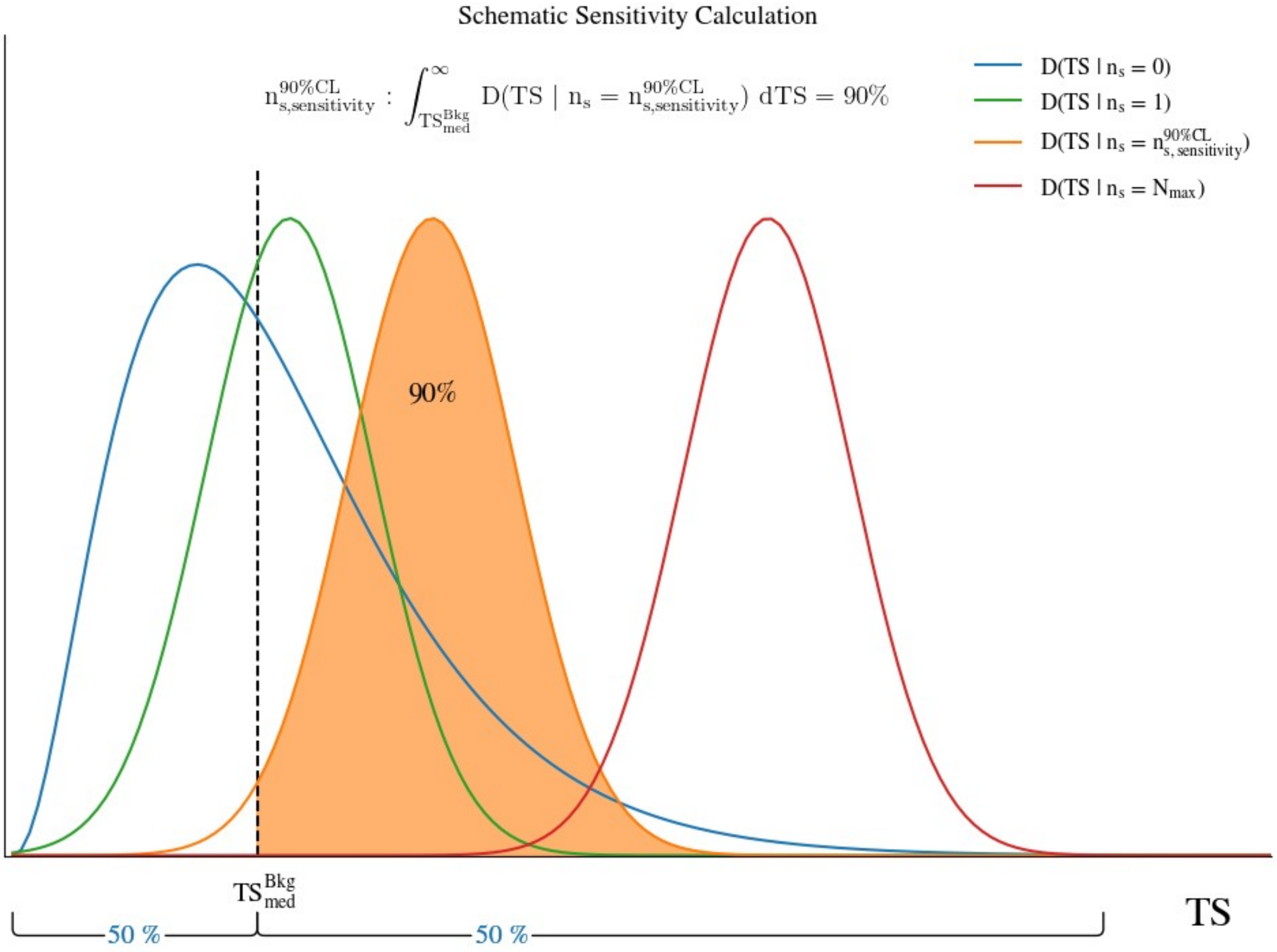
Analysis

$$\mathcal{L}(n_{\text{sig}}) = e^{-(n_{\text{sig}}+n_{\text{bkg}})} \prod_{i}^{N} \left[n_{\text{sig}} \cdot \mathcal{S}(\Psi_{\odot,i}, \beta_i, E_i) + n_{\text{bkg}} \cdot \mathcal{B}(\Psi_{\odot,i}, \beta_i, E_i) \right]$$

- Likelihood ratio test.

$$\text{TS} = \log_{10}\left(\frac{\mathcal{L}(\hat{n}_{\text{sig}})}{\mathcal{L}(0)}\right)$$

- Natural statistical fluctuations and 15% uncertainty in the number of detected events are included.
- Sensitivity computation.
- 90% CL upper limit computation.



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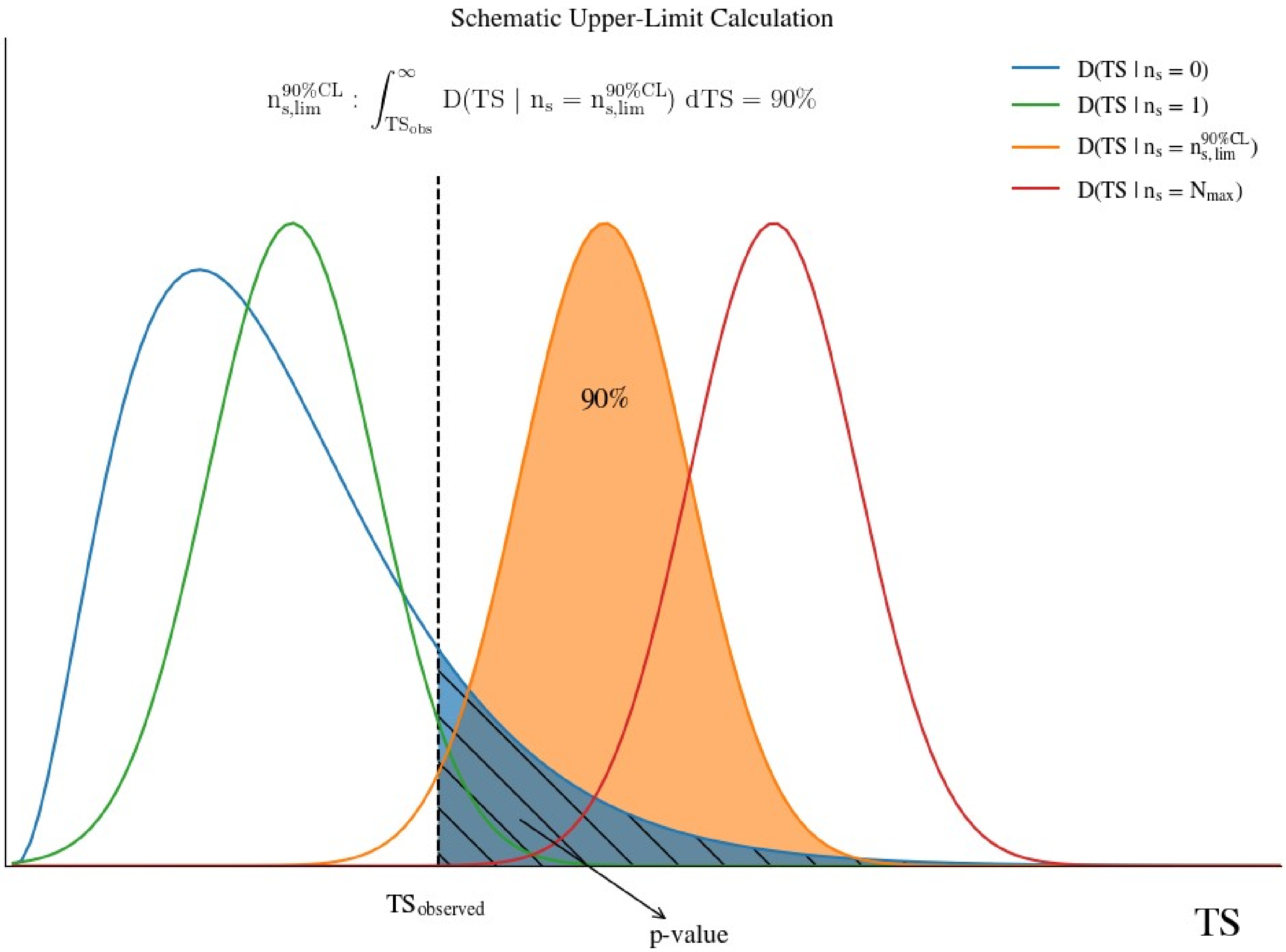
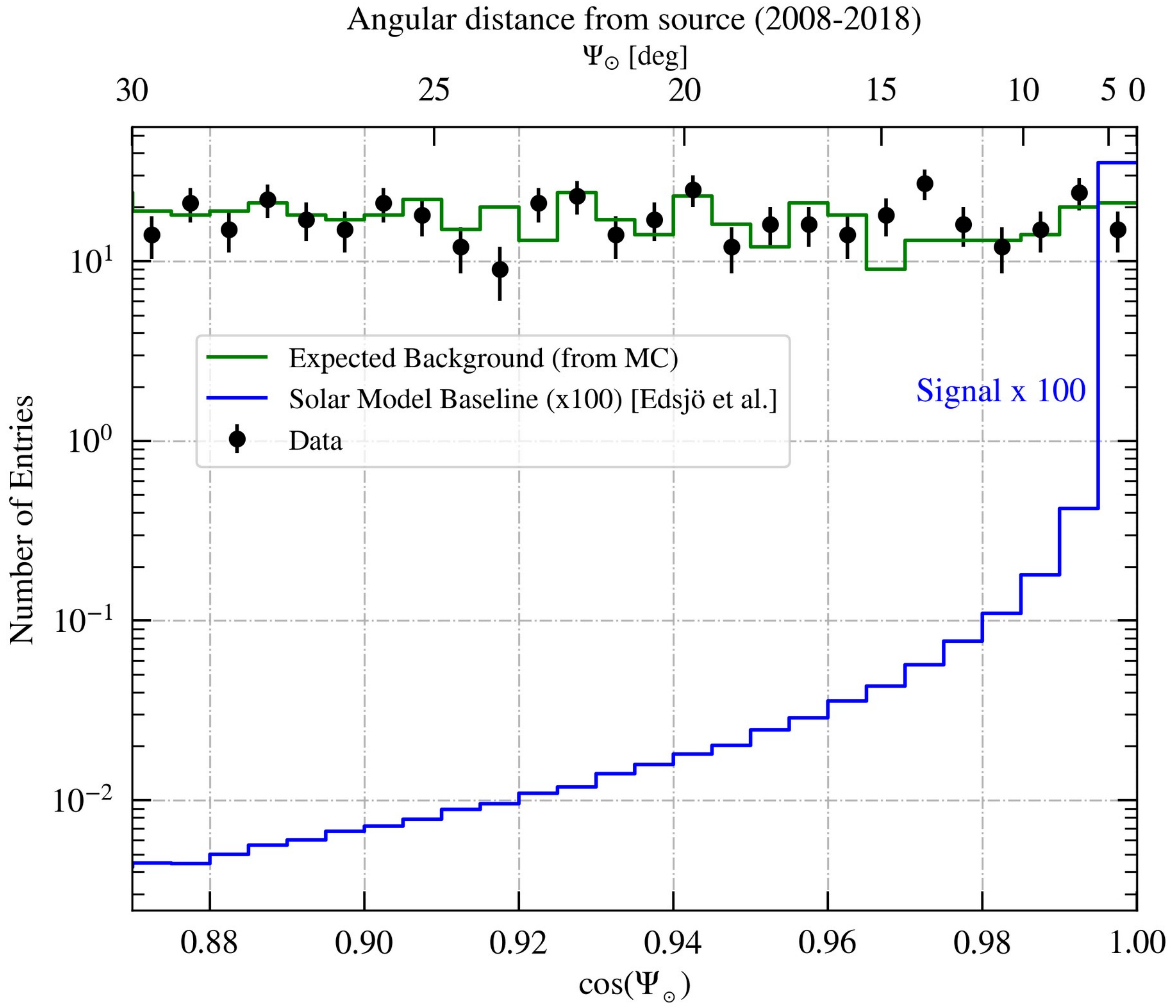


Analysis



- Event distribution as a function of the angular distance around the source.
- Expected signal magnified for comparison (blue histogram).
- Expected background (green).
- Data (black points).

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Analysis

- The flux limit is computed as:

$$\frac{d\phi^{90\%CL}_{\nu_\mu+\bar{\nu}_\mu}(E)}{dE} = \frac{\bar{\mu}^{90\%CL}_{sg}}{n^{\text{theor}}_{sg}} \cdot \frac{d\phi^{\text{theor}}_{\nu_\mu+\bar{\nu}_\mu}(E)}{dE}$$

Where:

$$n^{\text{theor}}_{sg} = T_{\text{live}} \int \sum_{l \in \nu_\mu, \bar{\nu}_\mu} \left(\frac{d\phi^{\text{theor}}_l(E')}{dE} A^l_{\text{eff}}(E') \right) dE'$$

- Is the expected number of signal events for the considered lifetime (3022 days).

- Unblinded results:
  - **$\mu_{90} = 3.15 \rightarrow C_{90} \approx 8.6$**
  - p-value = 0.41
  - Significance = 0.22σ